

# A Frequency Selective Surface Based Focal Plane Receiver for the OLIMPO Balloon-Borne Telescope

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**Abstract**—We describe here a focal plane array of Cold-Electron Bolometer (CEB) detectors integrated in a frequency selective surface (FSS) for the 350 GHz detection band of the OLIMPO balloon-borne telescope. In our architecture, the two-terminal CEB has been integrated in the periodic unit cell of the FSS structure and is impedance matched to the embedding impedance seen by it and provides a resonant interaction with the incident submillimeter radiation. The detector array has been designed to operate in background noise limited condition for incident powers of 20 pW to 80 pW, making it possible to use the same pixel in both photometric and spectrometric configurations. We present high-frequency and dc simulations of our system, together with fabrication details. The frequency response of the FSS array, optical response measurements with hot/cold load in front of optical window, and with variable-temperature blackbody source inside cryostat are presented. A comparison of the optical response to the CEB model and estimations of noise equivalent power (NEP) is also presented.

**Index Terms**—Cold-electron bolometer (CEB), frequency selective surface (FSS), OLIMPO.

## I. INTRODUCTION

A frequency selective surface (FSS) is an artificial arrangement of simple metallic repeating patterns that provide a resonant interaction with incident electromagnetic waves. The interaction manifests as bandpass, bandstop, reflecting, or absorbing behaviors [1]. A large number of simple repeating motifs have been used over the years in the form of crosses, squares, circles, and hexagons, arranged either on planar substrate or as free-standing metallic membranes that are used for quasi-optical mesh filters [2]. The bandpass filter is usually a mesh design with narrow bandwidth and is a candidate for integrating

bolometric detectors. A number of attempts have been made to use FSS-based filters fabricated on membranes with the working principle that they couple to incident power in a given bandwidth heating the wideband sensitive membrane [3], [4] and the increase in the membrane temperature is detected by a Transition Edge Sensor. Even multicolor radiometers have been fabricated and tested based on this idea [4], [5].

The authors of this paper have been involved with the cold-electron bolometer (CEB) detector [6] which is a two-terminal superconducting detector that uses two superconductor–insulator–normal metal tunnel junctions to couple power to a normal metal nanoabsorber. The junctions are also used to measure the temperature change of the nanoabsorber. The most important feature of the CEB architecture is that the same two junctions cool of the electron system of the absorber using the bias current which provides negative electrothermal feedback. In this paper, we describe detector architecture for the OLIMPO experiment [7]. In our planar pixel, we use the FSS approach for quasi-optical coupling to incoming signal power from space and detection is provided using the CEB detector. The CEB is a two-terminal detector, and we have managed to arrange it into the repeating FSS motif, e.g., annular shape or meander shape, and match the impedance of the CEB to the embedding impedance seen by the CEB in such a way that RF currents flow through the CEB heating the electron system of the normal metal and creating a response signal. Such an integration of the planar imaging array with the FSS results in absorbing of the incoming radiation by the detector only in a narrow frequency band. The readout for this response uses the same FSS unit cells connected in series and parallel [6], [8] such that the noise equivalent power (NEP) of the system together with the readout amplifiers is minimized.

The OLIMPO mission is a long-duration balloon-borne experiment with a  $f/3.44$  telescope and with detectors in four frequency bands cooled down to approximately 300 mK. The goal of the experiment is photometric and spectrometric observations of the Sunyaev-Zel'dovich effect from galaxy clusters. The Band 3 of the telescope has 23 pixels in the focal plane. Commercial quasi-optical mesh bandpass filters will be used to define the bandwidth from 330 to 365 GHz. The detectors are expected to have a NEP to be lower than photon noise. The estimated background power loads are 38 and 66 pW in the photometric and the spectrometric configuration, respectively, which sets the background noise level. The readout for CEBs will be cold JFET amplifiers similar to [9] and [10].

Manuscript received April 27, 2014; revised July 31, 2014; accepted September 26, 2014. This work was supported by the Swedish National Space Board.

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Digital Object Identifier 10.1109/TTHZ.2014.2362010

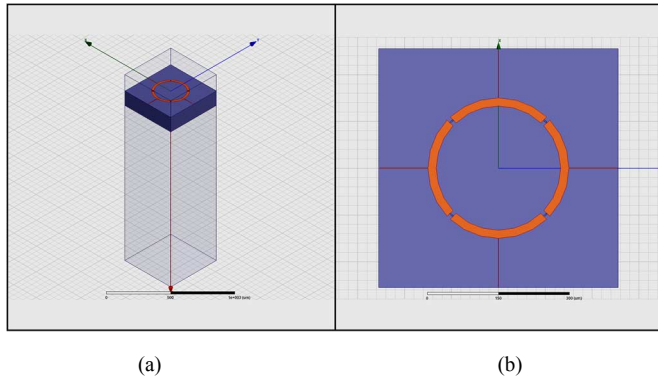


Fig. 1. (a) Unit cell of the FSS. (b) Close-up view of the annular mesh element. The Floquet excitation is arranged at the bottom the unit cell. The gray solid is the silicon substrate with the annular shape visible on it. The top is capped by the backshort. Periodic boundaries are arranged on each parallel face to simulate an infinite array and the phase difference between the periods is used to calculate the angle-dependent absorption of the FSS. The four bolometers (as four lumped elements) are arranged between the gaps of the annular element. The thin lines facilitate dc connection to neighboring pixels but are also an integral part of the FSS, being included in the RF simulations.

## II. FOCAL PLANE ARRAY OF DETECTORS

### A. FSS With Integrated CEB Detectors

We describe here the unit cell of the FSS structure in which the CEB detectors have been integrated. It is an annular shape with extended lines on all four sides. One opposite pair of lines terminate at a short distance from the unit cell edge such that the other pair of lines can facilitate dc connections to neighboring cells in only one direction (X or Y). This pattern is designed to be fabricated on a silicon substrate of calculated thickness backed with a backshort. The radiation is coupled from the silicon side similar to traditional lens antennas. The FSS surface together with the substrate and backshort form a resonant interference filter that couples to submillimeter radiation, where the backshort and the substrate tune out the reactance of the FSS layer. This structure has been simulated in the commercial 3-D electromagnetic simulation package HFSS using periodic boundary conditions and illuminated by a plane-wave excitation using a Floquet port. The unit cell structure is shown in Fig. 1 with four integrated detectors in the arms of the annular element. The return loss ( $S_{11}$ ) plot is shown in Fig. 2. In our model, the surface loss of the metal parts made of golden films was taken into account. Representing the four CEB detectors each by a lumped port with the same impedance, we have found that the amount of power dissipated in the gold electrodes is negligible compared with the power absorbed in the detectors. So, the relation  $\sum_{n=1}^4 S_{1n}^2 + S_{11}^2 = 1$  is satisfied, where  $S_{1n}$  is the power coupled to the  $n$ th detector. Thus,  $S_{11}$  can be used to describe the optical coupling to the detectors while simulating the complete pixel with hundreds of lumped elements. The best case coupling has been calculated for a unit cell with a silicon substrate thickness of  $141 \mu\text{m}$ , backshort distance of  $160 \mu\text{m}$ , and a unit cell pitch of  $510 \mu\text{m}$ . The design has been optimized such that the best coupling is obtained for a real part of detector impedance of  $50 \Omega$  and an imaginary part of impedance corresponding to a capacitance of  $50 \text{ fF}$ . These are very common parameters in CEB fabrication.

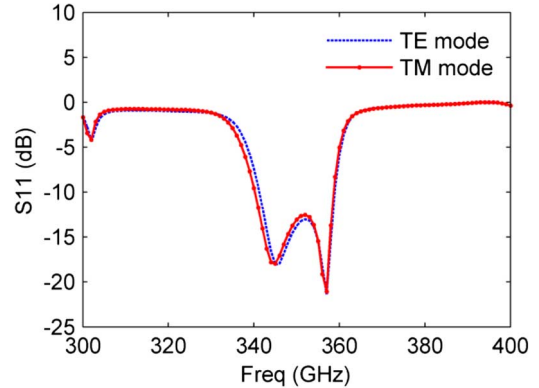


Fig. 2.  $S_{11}$  plot of the FSS unit cell for two orthogonal modes of the Floquet port.

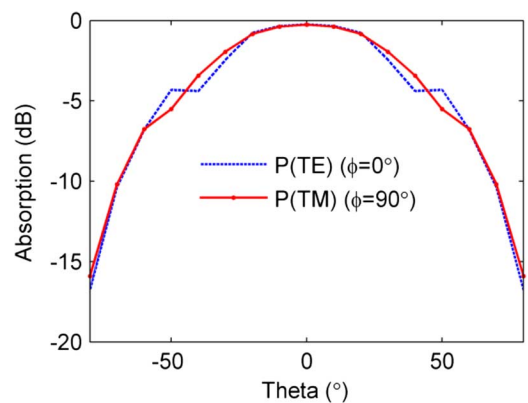


Fig. 3. Angle-dependent power absorption of the FSS for the two orthogonal (TE and TM) modes.

### B. Cold Stop

Planar incoherent arrays of detectors can couple to incident radiation over a large solid angle, and the advantages and disadvantages have been analyzed in [11]. In the current design, we have made preliminary calculations of the absorption of an arbitrarily polarised wave incident on the unit cell at different angles. Once the return losses  $S_{11}$  (TE) and  $S_{11}$  (TM) for TE and TM modes are known from numerical simulations, we can estimate the angle-dependent absorption of a finite pixel size by the formula  $P(\text{TE, TM}) = (1 - S_{11}^2(\text{TE, TM})) \cos \theta$ . The factor  $\cos \theta$  accounts for the different pitch to the incoming radiation incident at angle  $\theta$  relative to the focal plane normal vector. The result is shown in Fig. 3. Here we neglected the edge effects and obtained a beam width of  $60^\circ$  at a level of 3 dB. It is a narrower beam compared with that defined simply by the  $\cos \theta$  function. The reason is that, when the incident wave is not normal to the array plane ( $\theta > 0$ ), the excitation signal has a phase delay across the unit cell that can be represented as an additional surface reactance. This reactance increases with  $\theta$  and reduces the impedance match.

For a pixel mounted in the cryostat of the OLIMPO telescope, the “sidelobes” should not illuminate any surface of the cryostat that could potentially radiate some power in the bandwidth of the filters. Due to thermal budget restrictions, a black absorbing surface nearest to the detectors could not be considered. Instead,

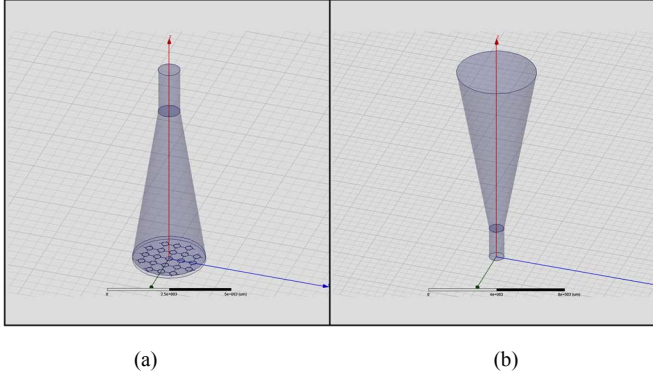


Fig. 4. (a) Model in HFSS of the lower part of the cold stop with CEB integrated FSS. (b) Model of the upper part of the cold stop.

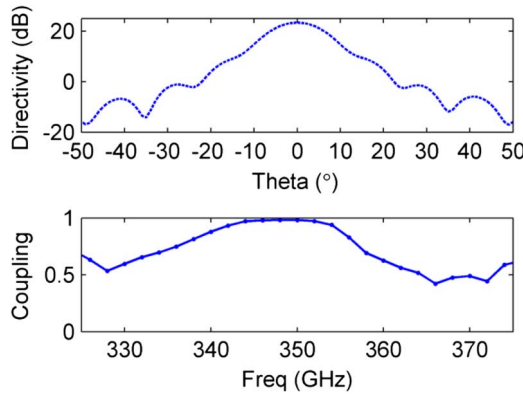


Fig. 5. Top: simulated directivity of the horn facing the telescope optics. Bottom: simulation of the optical coupling of single pixel with cold stop.

a new design of an array of horn based cold apertures was developed to limit the angle of visibility of the pixel. These apertures would be mounted on the 300 mK stage of the cryostat atop the pixels. This aperture is shown in Fig. 4 and consists of two back-to-back horns connected together with a 2 mm long section of a circular waveguide of diameter 0.9 mm that is slightly overmoded (five modes) at the frequencies of interest. The top part of the cold stop can be imagined as a traditional horn illuminating the telescope and was simulated as such in the commercial package HFSS. The bottom part of the cold stop can be simulated by assuming a waveguide port at the circular waveguide end exciting all possible waveguide modes and propagating the power in the direction of the FSS pixel. The return loss ( $S_{11}$ ) for each mode excited in the waveguide is a good indicator of the optical coupling of the pixel.

We have simulated the cold stop+FSS in HFSS. The simulations of the top part of the cold stop indicate that it can illuminate the telescope optics reasonably well with an on-axis directivity of 23.4 dB and half power beam width of 12°. This directivity pattern is shown in Fig. 5 (top). The simulation of the bottom part of the cold stop with the FSS also shows promising results. For the full pixel of size circa 3 mm × 3 mm, the optical coupling is better than 50% over the whole band. These results are shown in Fig. 5.

### C. DC Design

The dc design for the integrated planar array follows from [12] and [13]. Depending on the number of unit cells required to effectively cover the pixel area under the cold stop, the bolometers can be connected in series/parallel to match the noise impedance of the cold JFET amplifiers. In the current design, the cold stop limits the area of the pixel to a circle with a diameter of 3 mm where 21 unit cells of the annular shape can be fitted, each containing four bolometers, giving a total of 84 bolometers. The performance of the array can be simulated by modelling the performance of 1 CEB and scaling the responsivity to  $M$  array elements each experiencing  $M$  times less incident power. This analysis can be done using the heat balance equation [13]–[15] that equates power incoming into the normal metal absorber of the CEB and the power leaving the absorber due to bias current at quasi steady state:

$$2P_N + 2\beta P_S - \Sigma\Lambda(T_e^5 - T_{ph}^5) + I^2 R_a + \frac{P_0 + \delta P(t)}{M} = 0 \quad (1)$$

where  $P_N$  is the power deposited in the normal metal due to bias current,  $P_S$  is the power deposited in the superconducting electrode,  $\beta$  is a parameter [16] that refers to the fraction of power deposited in the superconducting electrode that returns back to the normal absorber,  $\Sigma\Lambda(T_e^5 - T_{ph}^5)$  is the heat flow from electron to phonon subsystems in the absorber,  $\Sigma$  is the electron-phonon coupling constant which is material-dependent,  $\Lambda$  is the volume of the normal absorber,  $T_e$  is the electron temperature of the absorber,  $T_{ph}$  is the bath temperature.  $I$  is the bias current,  $R_a$  is the resistance of the absorber,  $P_0$  is the background power on the pixel in steady state, and  $\delta P(t)$  is the incident signal power. Assuming that the signal power is a very tiny variation on top of the background power (this is especially true for balloon-borne experiments with very large atmospheric power loads), the pixel dc parameters can be estimated using the steady-state part of the heat balance equation.

The current through a SIN junction is dependent on the number of occupied and unoccupied states in each of the metals and can be expressed as [14]

$$I = \frac{1}{eR_N} \int_{-\infty}^{\infty} v(E) (f_N(E - eV) - f_s(E)) dE \quad (2)$$

where  $e$  is the electron charge,  $R_N$  is the normal resistance of the junction and  $V$  is the voltage across the same junction.  $v(E)$  is the Dynes density of states [17] for the superconductor and  $f_N(E)$  is the Fermi–Dirac distribution for the normal metal which depends not only on the energy but also on the temperature of the electron subsystem in the normal metal. This dependence of temperature is what we correlate to absorbed power in the normal metal.  $f_s(E)$  is the Fermi–Dirac distribution for the superconductor, and, in general, it will be different from  $f_N(E)$  since they can have different electron temperatures.

The tunneling electrons in the SIN junction will carry with them some power; this power  $P_N$  can be calculated by removing one  $e$  in (2) and replace it with  $eV - E$  which is the energy

deposited in the normal metal from one tunneling event. Doing this, we obtain [14]

$$P_N = \frac{1}{e^2 R_N} \int_{-\infty}^{\infty} (eV - E)v(E) (f_N(E - eV) - f_S(E)) dE \quad (3)$$

It is possible for  $P_N$  to be negative and this is the effect we use to cool the normal metal, giving the detector its name.

From energy conservation, we conclude that the power dissipated in the superconductor is

$$P_S = IV - P_N. \quad (4)$$

The responsivity of the CEB in current-bias mode is given by

$$S_V = -2 \frac{\frac{\delta I}{\delta T} / \frac{\delta I}{\delta V}}{G_{e-ph} + 2G_{\text{sin}}} \quad (5)$$

where  $G_{e-ph}$  is the heat conductance due to electron–phonon coupling and  $G_{\text{SIN}}$  is the heat conductance of one tunnel junction [15].  $AI/AV$  and  $\delta I/\delta V$  are the partial derivatives of the current with respect to temperature and voltage.

The total NEP of the system can be divided into three parts, the SIN-junction itself, the electron–phonon interaction, and the amplifier noise. The squares of the three parts can be calculated as

$$\text{NEP}_{\text{SIN}}^2 = \delta P^2 - 2 \frac{\delta P \delta I}{\delta V S_V} + \frac{\delta I^2}{\left(\frac{\delta I}{\delta V} S_V\right)^2} \quad (6)$$

$$\text{NEP}_{e-ph}^2 = 10k_B \Sigma \Lambda (T_e^6 + T_{ph}^6) \quad (7)$$

$$\text{NEP}_{\text{amp}}^2 = \frac{v_n^2 + \frac{N}{W} \left( I_n^2 \left( \frac{\delta v}{\delta I} + R_a \right) \right)^2}{\left( \frac{S_v}{W} \right)^2}. \quad (8)$$

In the SIN junction, there exist three different sources of noise;  $\delta P^2$  is the heat flow noise,  $\delta I^2$  is the fluctuations in the current or shot noise and since the current and the heat flow are correlated there also exists a correlation term  $\delta P \delta I$  [15]. The amplifier contributes two noise sources,  $V_n$  the voltage noise and  $I_n$  the current noise. Assuming  $N$  is the number of bolometers in series and  $W$  is the number of bolometers in parallel, giving a total of  $N * W$  bolometers, the total NEP is then given by

$$\text{NEP}_{\text{total}}^2 = N * W * \text{NEP}_{\text{SIN}}^2 + N * W * \text{NEP}_{e-ph}^2 + \text{NEP}_{\text{amp}}^2. \quad (9)$$

The most optimized combination of these bolometers taking into account the geometry and lowest possible NEP over the full range of incident powers (20 to 80 pW, with a margin on both sides of OLIMPO estimations) is 42 bolometers in series and two in parallel. This corresponds to 21 unit cells of the FSS.

The NEP plots for two different power loads calculated from (1)–(6) are shown in Fig. 6. We notice from the plots that the pixel with  $42 \times 2$  bolometers performs well with NEP of the device less than the photon noise over a range of powers from 20 to 80 pW. Thus, the same pixel can be used for both the photometric and spectrometric configurations, when, in the spectrometer configuration, the background power is a lot larger owing to emissions of uncooled spectrometer mirrors.

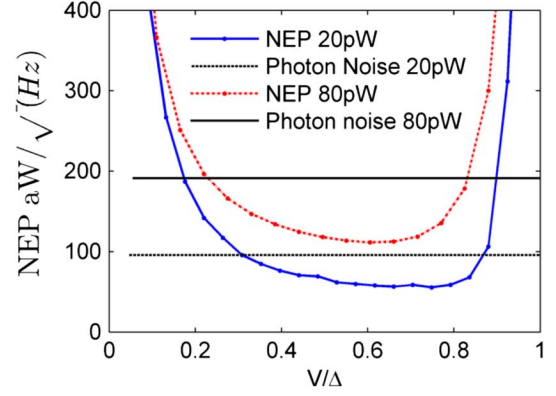


Fig. 6. Contributions of noise to the pixel design with the voltage scale normalized to one junction in the array for power loads of 20 and 80 pW. This plot has been calculated for  $R_a = 150\Omega$ ,  $R_n = 1k\Omega$ , Volume of absorber  $\Lambda = 0.023 \mu\text{m}^3$ ,  $\beta = 0.1$ ,  $\Sigma = 1.5 \text{ nW}/(K^5 \mu\text{m}^3)$ . Amplifier voltage noise and current noise  $I_n = 3fA/\sqrt{Hz}$ .

#### D. Fabrication

The fabrication process is similar to the one described in [18]. The pixel is fabricated using multiple lithography layers. The initial layer consists of alignment marks and defines the pixel size. The next layer is used to expose and evaporate contact pads and the FSS structure. Silicon machining using deep Silicon etch process constitutes the third layer. The Silicon substrate is thinned down to the required thickness using the Bosch process. The Bosch process used for deep silicon etching uses the STS ICP plasma etcher with  $\text{SF}_6$  gas in the etch step and  $\text{C}_4\text{F}_8$  gas in the passivation step providing an anisotropic etch of the silicon. The process is stopped at the required depth with an etch error typically less than  $2 \mu\text{m}$ . The final layer is the exposure of the CEB devices in an electron-beam lithography machine. The size of each of the two bolometer SIN tunnel junctions is  $2 \mu\text{m} \times 0.4 \mu\text{m}$  connected with the nanoabsorber being  $2 \mu\text{m}$  long,  $100 \text{ nm}$  wide, and  $14 \text{ nm}$  thick. After exposure and development of the e-beam resist, the bolometer layers are evaporated. The normal metal absorber is a layer of aluminum whose superconductivity is suppressed by a thin layer of ferromagnetic impurities (Fe). The tunnel barrier of the SIN tunnel junction is formed by oxidation of the normal metal at 10 mbar oxygen pressure for 5 min. The superconducting electrodes are then evaporated at two angles ( $+45^\circ$  and  $-45^\circ$ ) to a thickness of  $70 \text{ nm}$  each. Finally, all unwanted metal and resist is lifted off using warm acetone and a gentle ultrasonic agitation.

#### E. Prototype Testing

The first prototype pixel was fabricated and characterized at low temperatures. The test consisted of measuring voltage response of the device in front of open cryostat windows with room temperature and liquid nitrogen loading. The incoming radiation was filtered with quasioptical bandpass filters along with Fluorogold to decrease infrared power. The optical response to  $300 \text{ K}/77 \text{ K}$  loading is shown in Fig. 7. The  $I$ – $V$  curves of the device are shown on the left axis; the curve with lower linearity corresponds to  $300 \text{ K}$  radiation. The optical response is plotted on the right axis as the difference in voltage of the two  $I$ – $V$  curves. This test highlights the optical sensitivity of the device.



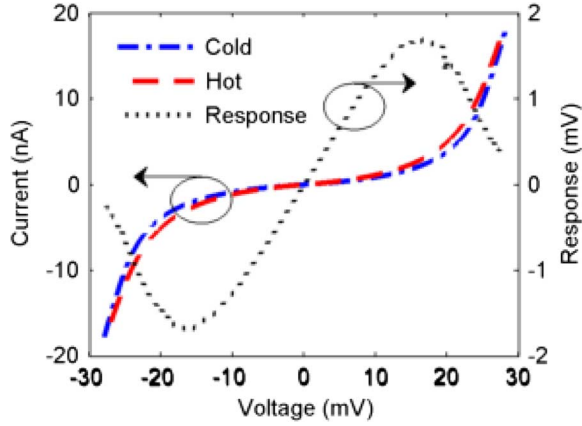


Fig. 7. Optical response with 300 K and 77 K loading.

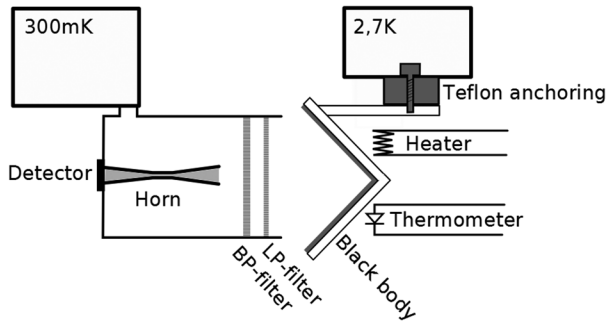


Fig. 8. Prototype measurement setup.

The power incident on the device with the optical load at 300 K is more than 550 pW and at 77 K is more than 150 pW. Assuming the optical coupling from Fig. 5, the device is operating beyond its designed power load and is overheated. This causes the responsivity to be far lower than at design power loads; here estimated about  $4.5 \times 10^6 (V/W)$ .

For testing the device at lower power loads, a test setup with a cold blackbody source shown in Fig. 8 was used. The prototype was connected to the 300 mK stage of a cryostat. The sample holder had blackened inner surface to reduce spurious reflections and low-pass and bandpass filters to define the correct optical band. The RF source was a cold black body with a heater and thermometer connected to the 2.7 K stage of the cryostat using a dielectric anchor to prevent heat leaks to the cryostat shields. The incident power on the pixel could be varied by changing the black body temperature between 2.7 and 6 K. The incident power could be estimated using Planck's law. Current-voltage curves and voltage-power curves of the CEB detectors were recorded by changing the blackbody temperature. Measurement of the response of the pixel is shown in Fig. 9. The estimated responsivity is also shown in the figure; this is calculated using estimated incident power on the device. The responsivity decreases with increasing incident power which is due to increasing electron temperature of the absorber. In this calculation, the black body emissivity was assumed to be 0.9, and the blackened inner surface of the sample holder is assumed to attenuate any stray radiation falling on it. The fluorogold cover to

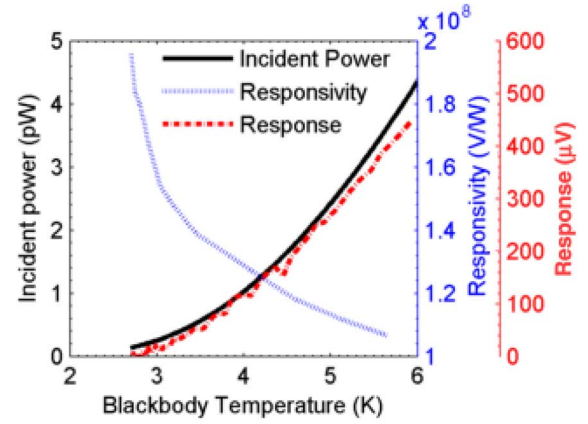


Fig. 9. Incident power, optical response, and estimated responsivity of the pixel.

the sample holder was assumed to have a transmission of unity in the bandwidth of interest.

Another way to estimate the responsivity of the device is by fitting the optical response data to the bolometer model presented in (1)–(5). An  $I-V$  curve of the device is taken with the black body kept at a known temperature. These data are then fitted to the model described earlier and responsivity is calculated. For small changes in radiated power from the black body, this responsivity curve can very reasonably describe the change in absorbed power in the bolometer array. Thus, the value of absorbed power can be estimated. Fig. 10 shows data fits to optical response taken at blackbody temperatures of 2.7 K, 3.6 K, and 4.5 K. The two response curves are generated by subtracting individual  $I-V$  curves at one blackbody temperature from the lowest one. The data fits indicate that approximately 0.2 pW of power was absorbed at 3.6 K and approximately 1 pW power was absorbed by the array at 4.5 K blackbody temperature. Compared with Fig. 9, the estimated absorbed power is lower by about a factor of 2, and the estimated responsivity is correspondingly higher. The difference can be attributed to over estimation of emissivity of the black body, under estimation of attenuation of one of the elements in the optical path, difference in the frequency response of the array with the cold stop compared to simulations, or a combination of these instances.

Using the above estimated responsivity, one can also estimate the electrical NEP by dividing the measured value of noise (at 115 Hz) by the responsivity. This is presented in Fig. 11. The background power is estimated to be 6 pW from the data fit, and the corresponding photon noise level is indicated. Using the (6)–(9), the calculated NEP for the actual parameters of the fabricated sample is also indicated. The current and voltage noise values are referred to the commercial JFET amplifier AD743 beyond its  $1/f$  corner. This amplifier was used for noise measurement in the actual experiment.

The estimated NEP is approximately  $1.5 \times 10^{-16} (W/\sqrt{Hz})$  at optimum bias point and is approximately four times higher than the modeled one. At higher absorbed power, the dynamic resistance of the array will be lower, and the impact of the current noise of the JFET amplifier will be reduced. The measurement showed low frequency interferences that were the likely

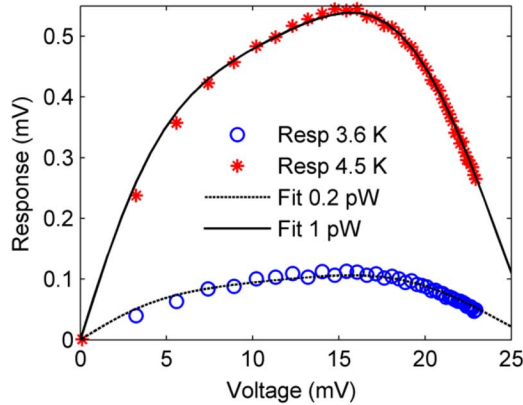


Fig. 10. Optical response of the array at two black body temperatures and data fits for absorbed power from the CEB model. This fit has been calculated for  $R_a = 150 \Omega$ ,  $R_n = 2 \text{ k}\Omega$ , Volume of absorber  $\Lambda = 0.023 \mu\text{m}^3$ ,  $\beta = 0.03$ ,  $\Sigma = 1.1 \text{ nW}/(\text{K}^6 \mu\text{m}^3)$ .

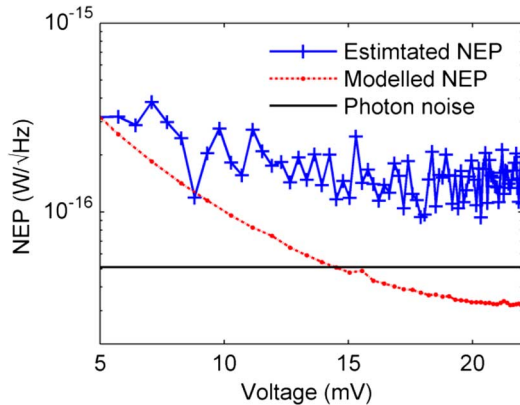


Fig. 11. Comparison of estimated NEP, modelled NEP and photon noise. Amplifier voltage noise  $V_n = 4nV/\sqrt{H_Z}$  and current noise  $I_n = 5fA/\sqrt{H_Z}$ .

source of the increased noise and of the 6 pW background power indicated earlier. One source of the low-frequency interference is the vibrations of the pulse tube cooler of the cryostat used in the experiment. It was not possible to switch it off to reduce its effects since it introduced temperature gradients during the course of the measurement.

To understand the frequency response of the FSS, an experiment was made where the FSS was illuminated with radiation from a backward wave oscillator (BWO) through a bandpass filter chain. Initially, the filter chain was characterized separately and then used in the measurements of the FSS. Given the variation of BWO emitted power versus the operating frequency, we normalized the measured voltage with the corresponding power reporting the normalized efficiency in the radiation detection. The results are shown in Fig. 12. The filter frequency response has a sufficiently large bandwidth to cover the FSS band. The FSS itself has a frequency response comparable to the simulations shown in Fig. 2.

### III. DISCUSSION

Compared with the use of traditional coupling approaches like the spiderweb bolometer or the slot antennae, the distributed FSS array is conducive to the use of two terminal CEB devices.

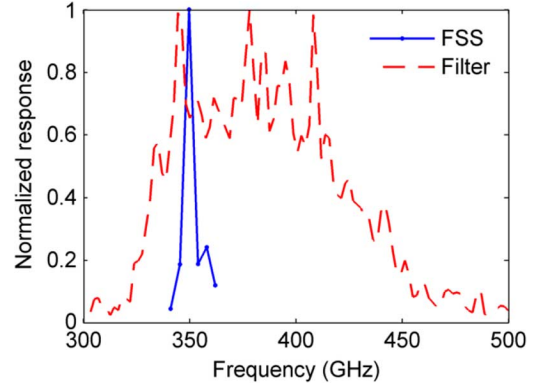


Fig. 12. BWO measurement of the frequency response of the FSS and the bandpass filter characteristics.

In the case of the spiderweb or the slot antennae, a single (or few) detector absorbs power that is coupled to the spiderweb or through the immersion dielectric lens that usually accompanies the slot antenna. In the case of the FSS array, incident power is distributed over a number of bolometers such that the array of CEBs acts as a distributed absorber along with a distributed coupling element (the FSS geometry). This distribution of power causes each bolometer to be “cooler” which improves the responsivity. In the case of OLIMPO, it was not possible to design a single CEB element that could handle the specified power, so the FSS concept was developed.

The CEB detectors could be used in large-scale arrays, which are needed for almost all future instruments. The device fabrication is not an issue since facilities and competence exists to fabricate arrays over the size of up to 6-in wafers (> hundred pixels). The most important consideration is the readout system. The multiplexing readout technology for medium/high-impedance sensors is relatively undeveloped. Analog multiplexers based on Si-Ge ASICs could solve the problem [19], by means of low-noise MOSFET switches operating at low temperature with low power dissipation, but this has not been tested with the arrays. Another solution is low-ohmic arrays matched to superconducting quantum interference device (SQUID) readout. For this goal, well-developed SQUID readout systems with multiplexing could be used [20]. For matching of relatively high-ohmic CEB with low-ohmic SQUID, a parallel array of CEBs was proposed for realization of background-limited operation [8]. If impedance matching of the CEB array and SQUID is not sufficient, an additional superconducting transformer could be effectively used for matching.

### IV. CONCLUSION

We have designed and fabricated a single pixel of a submillimeter focal plane receiver using FSS-based arrays integrated with the CEB detector. Optical response was measured with the cold stop designed for the OLIMPO project and responsivity of  $2 * 10^8 (\text{V}/\text{W})$  was estimated at low power loads and  $4.5 * 10^6 (\text{V}/\text{W})$  at high power loads. Electrical NEP for this array has been estimated to be  $1.5 * 10^{-16} (\text{W}/\sqrt{\text{Hz}})$  at low background power. These results of the prototype testing indicate very promising characteristics of the pixel. We hope to use

this approach for future submillimeter focal-plane-based pixel arrays.

#### ACKNOWLEDGMENT

The authors would like to thank L. Jönsson for assistance with the fabrication of the cold stop and the sample holder and A. Schillaci and G. D'Alessandro for assistance with the FTS measurements of the filter chain transmission.

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